

# Numerical simulations of the CME on 2010 April 8

Yingna Su<sup>1</sup>, Bernhard Kliem<sup>2</sup>, Adriaan van Ballegoijen<sup>1</sup>,  
and Edward Deluca<sup>1</sup>

<sup>1</sup>Harvard-Smithsonian Center for Astrophysics, Cambridge, MA, United States  
email: ynsu@head.cfa.harvard.edu

<sup>2</sup>Institute for Physics and Astronomy, University of Potsdam, Potsdam, Germany

**Abstract.** We present 3D zero-beta ideal MHD simulations of the solar flare/CME event that occurred in Active Region 11060 on 2010 April 8. The initial magnetic configurations of the two simulations are stable nonlinear force-free field and unstable magnetic field models constructed by Su *et al.* (2011) using the flux rope insertion method. The MHD simulations confirm that the stable model relaxes to a stable equilibrium, while the unstable model erupts as a CME. Comparisons between observations and MHD simulations of the CME are also presented.

**Keywords.** Sun: coronal mass ejections (CMEs), Sun: flares, Sun: magnetic fields.

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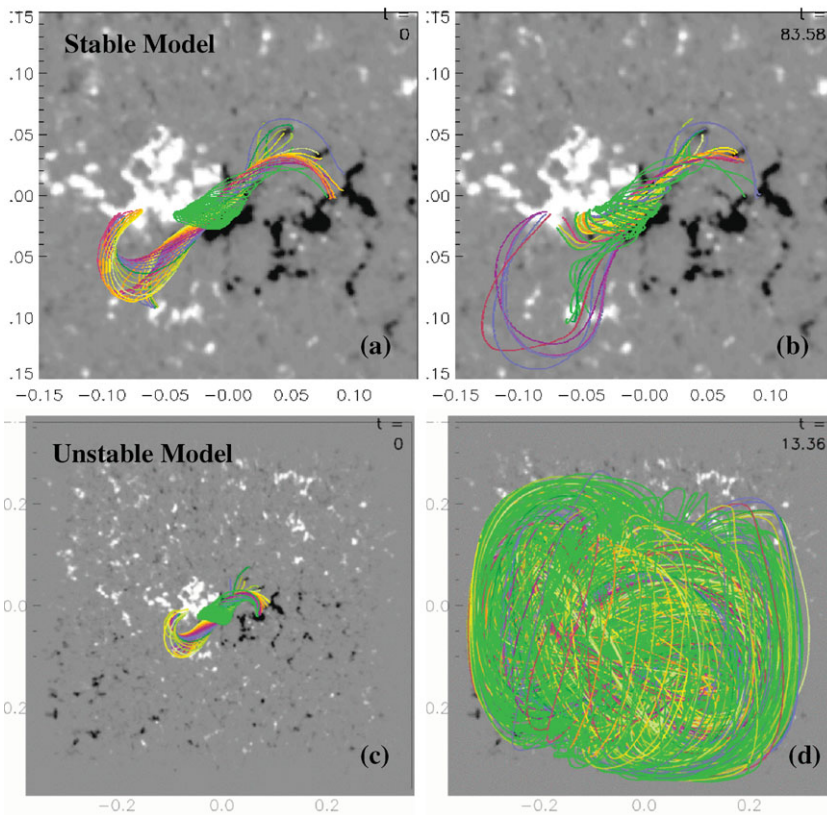
## 1. Introduction

A GOES B3.7 two-ribbon flare occurred in NOAA Active Region 11060 around 02:30 UT on 2010 April 8. The associated CME median speed is about 510 km s<sup>-1</sup>. In our earlier paper (Su *et al.* 2011), we constructed a series of static magnetic field models of the active region using the flux rope insertion method (van Ballegoijen 2004). We find that the axial flux of the flux rope ( $\Phi_{axi} = 4 \times 10^{20}$  Mx) in the best-fit pre-flare nonlinear force-free field model is very close to the threshold of instability ( $\Phi = 5 \times 10^{20}$  Mx). The unstable model ( $\Phi = 6 \times 10^{20}$  Mx) due to increase of axial flux closely matches the observations at the early stage of the eruption. The poloidal flux in these models remains the same, i.e.,  $F_{pol} = 10^{10}$  Mx cm<sup>-1</sup>.

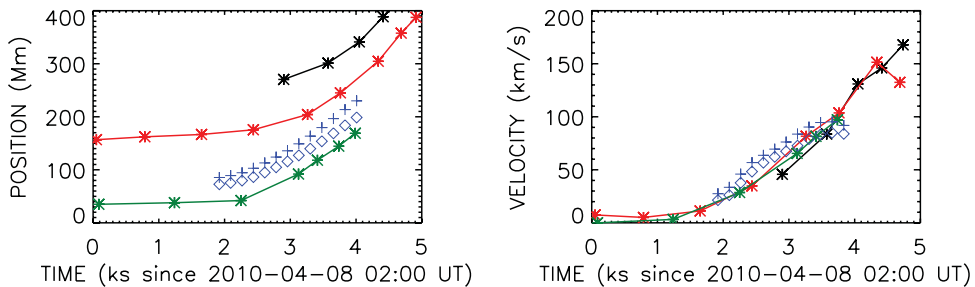
## 2. MHD Simulation versus Observations

In the present study, we perform 3D zero-beta MHD simulations (Kliem *et al.* 2004) using the aforementioned static magnetic field models as initial configurations. The magnetic field models are constructed in spherical geometry, while the MHD simulations are carried out in a cartesian geometry. The coordinate conversion is implemented by producing scaled down (scaling factor = 1/16.97) versions of the original magnetic models as inputs for the MHD simulations. The computation domain of the MHD simulation is .72×.72×1 solar radii. The top and side boundaries are closed. The model for the initial density is  $\rho(x, y, z, 0) = |B(x, y, z, 0)|^{1.5}$ .

Figure 1 shows selected field lines from the two MHD simulations using the stable (top row) and unstable (bottom row) models as initial conditions. As shown in Figure 1, the MHD simulations confirm that the stable model is really stable, while the unstable model erupts as a CME. STEREO\_A observations suggest that the rising path of the erupting filament is not radial, but rather inclined towards the equator by about 45°. This inclination is also identified in the MHD simulation of the unstable model. Figure 2 shows a good agreement between observations and simulations of the temporal evolution of the position and velocity of selected coronal loops.



**Figure 1.** Initial (left) and final (right) magnetic configurations of two MHD simulations of the event. The initial conditions of the two simulations are stable (top) and unstable (bottom) magnetic field models.



**Figure 2.** Comparisons between observations (asterisk) and MHD simulations (diamond and plus sign) of the temporal evolution of the position (left) and velocity (right) of selected coronal loops.

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## References

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