

Thermal characteristics of a B8.3 flare observed on July 04, 2009

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Abstract. We explore the temporal evolution of flare plasma parameters including temperature (T) - differential emission measure (DEM) relationship by analyzing high spectral and temporal cadence of X-ray emission in 1.6-8.0 keV energy band, recorded by SphinX (Polish) and Solar X-ray Spectrometer (SOXS; Indian) instruments, during a B8.3 flare which occurred on July 04, 2009. SphinX records X-ray emission in 1.2-15.0 keV energy band with the temporal and spectral cadence as good as 6 μ s and 0.4 keV, respectively. On the other hand, SOXS provides X-ray observations in 4-25 keV energy band with the temporal and spectral resolution of 3 s and 0.7 keV, respectively. We derive the thermal plasma parameters during impulsive phase of the flare employing well-established Withbroe-Sylwester DEM inversion algorithm.

Keywords. Sun: corona, Sun: flares, plasmas, Sun: X-rays, radiation mechanisms: thermal, techniques: spectroscopic.

1. Introduction

Thermal characteristics of solar flare plasma employing the multi-wavelength observations is of immense interest as it can shed light on the ongoing coupling processes in solar atmosphere. In particular, X-ray emission during a flare is the best probe of various thermal and non-thermal energy release processes (Brown (1971)). Generally, flare plasma parameters viz. temperature (T), emission measure (EM), etc. are derived by forward-fitting/inversion of the observed X-ray spectrum (Jain *et al.* (2011)). However, the spectroscopic inversion of X-ray emission is an ill-posed problem, leading to substantial uncertainties in the derived T and EM values (Craig & Brown (1976)). Moreover, several different DEM inversion techniques, with various functional dependence of DEM on T viz. power-law, single-gaussian etc., are used to interpret observed X-ray spectrum. Further, Withbroe-Sylwester (W-S) DEM inversion algorithm (Sylwester, Schrijver, & Mewe (1980), Kepa *et al.* (2008)) provides a more general scheme for such studies. Therefore, in this paper, we present the analysis of X-ray emission observed during a B8.3 flare occurred on July 04, 2009, the only event recorded in common with Solar X-ray Spectrometer (SOXS; Jain *et al.* (2005)) and Solar Photometer in X-rays (SphinX; Gburek *et al.* (2013)). Section 2 presents the observations while data analysis and results are given in Section 3. Section 4 presents the summary and conclusions.

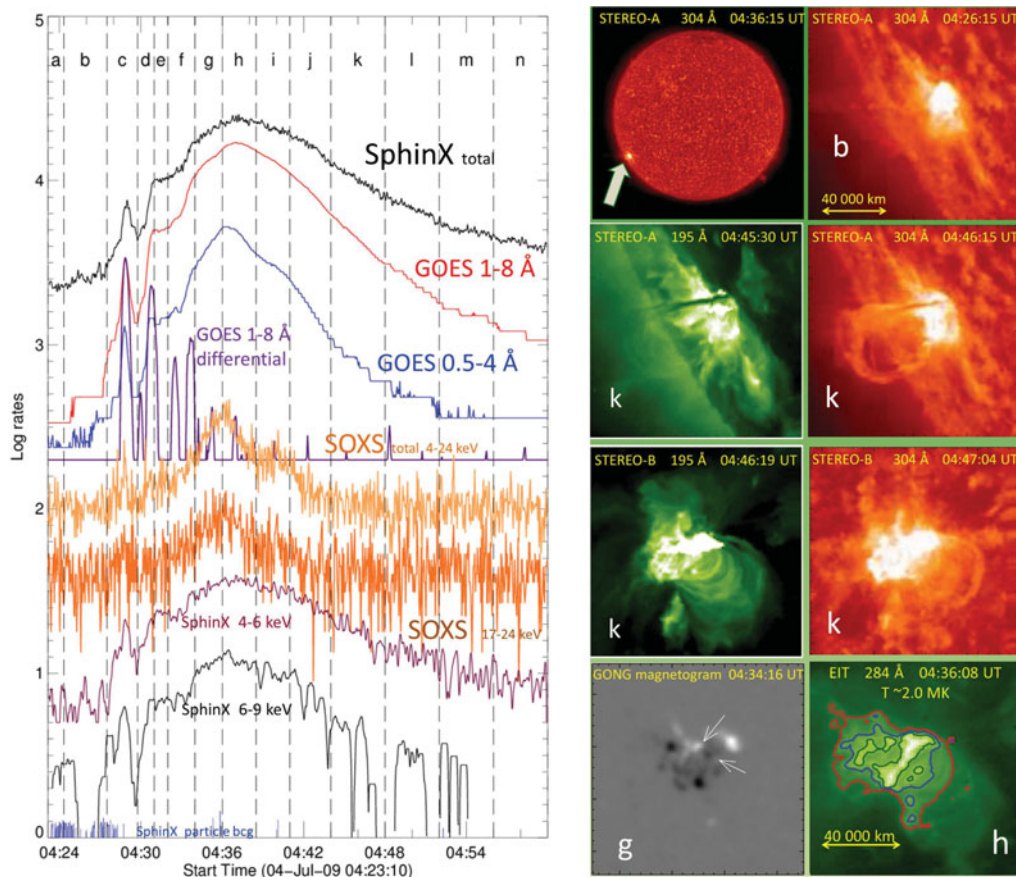


Figure 1. Left Panel: Temporal evolution of X-ray emission as recorded by SphinX, SOXS and *GOES* during SOL2009-07-04T04:37 (B8.3) flare. Dotted bars show the time intervals for which spectral analysis is undertaken. Right Panel: Multi-wavelength overview of the flare from *STEREO-A* and *STEREO-B* and EIT/*SOHO*. Activity areas are shown by arrows in the GONG Magnetogram (bottom).

2. Observations

We study a B8.3 flare event of July 04, 2009, which occurred in active region 11024. Thermal characteristics of the flare plasma are derived by analyzing X-ray spectra in 1.6–5.0 keV and 5.0–8.0 keV energy bands, recorded by SphinX and SOXS, respectively. Temporal evolution of X-ray emission during the flare as observed by SphinX and SOXS instruments as well as by *GOES* is shown in the left panel of the Fig. 1. Further, morphological evolution of the flaring region is studied from the EUV images obtained from *STEREO-A*, *B* and Extreme Ultraviolet Imager Telescope (EIT) onboard *SOHO* mission, as shown in the right panel of the Fig. 1.

3. Thermal characteristics of the flare plasma

We analyze the X-ray spectra, recorded during the flare, with the help of W-S *DEM* inversion method (Sylwester, Schrijver, & Mewe (1980)). This numerical method employs maximum likelihood approach in which a *DEM*–*T* distribution and hence corresponding theoretical spectrum is derived in an iterative manner with the aim to minimize its

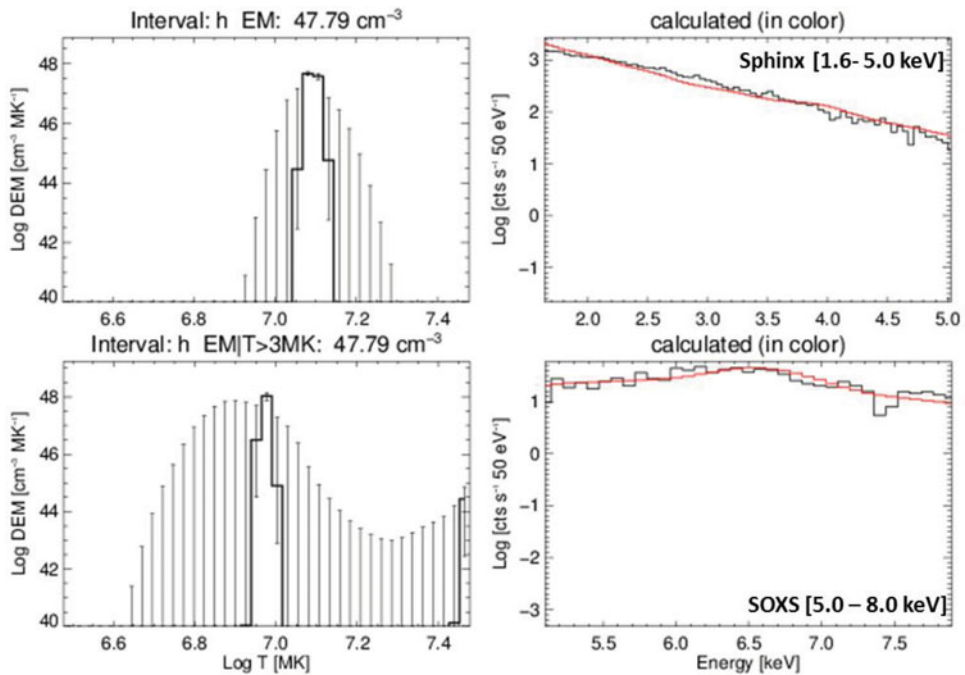


Figure 2. Top row: Best-fit $DEM-T$ relationship as well as the spectral-fit (drawn by red color) employing W-S inversion algorithm for the emission in 1.6-5.0 keV (plotted by black color), recorded by SphinX during 04:36:00-04:38:30 UT. Bottom row: Best-fit $DEM-T$ and fitted SOXS spectrum in 5.0 - 8.0 keV for the aforesaid time duration.

difference with the input observed spectrum (Kepa *et al.* (2008)). Coronal abundances are adopted from CHIANTI atomic database (Del Zanna *et al.* (2015)) while deriving the shape of theoretical spectra. As an input to this method, we have used fluxes recorded in the 73 energy bins (corresponding to the energy band 1.6-5.0 keV) and 35 energy bins (corresponding to the energy band 5.0-8.0 keV) by SphinX and SOXS instruments, respectively. The X-ray spectra are analyzed for various time duration as shown by dotted lines in the left panel of Fig. 1. Top row of the Fig. 2 shows the best-fit $DEM-T$ relation derived by analyzing X-ray spectrum in 1.6-5.0 keV (low-energy), observed by SphinX during the peak of the impulsive phase of the flare, 04:36:00-04:38:30 UT. Similarly, in the bottom panel, we present the best-fit $DEM-T$ curve and spectral-fit over the X-ray spectrum in 5.0-8.0 keV (high-energy), observed by SOXS during the aforesaid time.

From Fig. 2, it may be noted that the best-fit $DEM-T$ relation derived from SphinX observation suggests nearly isothermal nature of the DEM , with the peak at temperature (T_p) \sim 13 MK. Similarly, the best-fit DEM to the SOXS spectrum in 5.0-8.0 keV energy band for the same time interval suggests isothermal nature in the form of single gaussian function dependence on T , however, at $T_p = 9.8$ MK. Next, thermal energy are derived from the best-fit $DEM-T$ curve of SphinX and SOXS observations and estimated to be 5.1 and 3.6×10^{29} ergs, respectively. We employ the volume estimated from EIT/SOHO EUV wavelength images as shown in Fig. 1 for the calculation of thermal energies.

4. Summary and Conclusions

We study the thermal characteristics of the plasma during SOL2009-07-04T04:37 (B8.3) flare by analyzing its X-ray spectrum in various energy bands, as obtained by

SphinX and SOXS instruments. We summarize the preliminary findings of this study as follows:

(a) Emission-measure is found to be of isothermal nature during the peak of the impulsive phase of the flare.

(b) Thermal energy and the temperature estimated by analyzing low-energy (from SphinX) and high-energy (from SOXS) bands within SXR spectrum result in different peak temperature as well as thermal energy.

In the next step, we have made a detailed investigation of thermal characteristics as well as the evolution of *DEM-T* relationship in various phases of the flare by combining the observations from SphinX and SOXS instruments. The corresponding paper is under review in the ApJ Main Journal (Awasthi *et al.* (2016)).

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