

Final Fates of Massive Stars

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Abstract. Massive stars are thought to play important roles in the early evolution of the Universe. In this paper, we first classify the final fates of massive stars into 7 cases according to their mass ranges. These variations of the final fate may correspond to the observed large diversities of supernova properties, such as extremely faint and extremely luminous (superluminous) supernovae, and the extremely energetic hypernovae. We then focus on the properties of the peculiar superluminous Type Ic supernova 1999as. We examine radioactive decay models, magnetar models, and circumstellar interaction models for the light curve of SN 1999as. We find that these models are not quite successful, and thus it is crucially important to improve these models to clarify the final fates of massive stars.

Keywords. supernova, hypernova, superluminous, nucleosynthesis

1. Final Fates of Massive Stars

The final stages of massive star evolution, supernova properties, and their chemical yields depend on the progenitor's main-sequence masses M (e.g., Arnett 1996, Filippenko 1997, Nomoto *et al.* 2006, Smartt 2009). Here we call some specific supernovae (SNe) as follows. In terms of the kinetic explosion energy E , we use "Hypernovae" for such energetic SNe as $E_{51} = E/10^{51} \text{erg} > 10$. In terms of brightness, we use "Faint SNe" for low luminosity SNe, and "Superluminous SNe (SLSNe)" for SNe brighter than -21 mag at maximum (Quimby in this volume).

The following mass ranges are set by various types of criteria, based on some combinations of observations and models. But the criteria and critical masses are not quite systematic yet, and should still be regarded as working hypothesis (Nomoto *et al.* 2009, 2010).

(1) $M_{\text{up}} - 10 M_{\odot}$ stars: **Faint supernovae:** These stars become electron capture SNe because their degenerate O+Ne+Mg cores collapse due to electron capture. $M_{\text{up}} \sim 9 \pm 0.5 M_{\odot}$ depending on the mass loss rate on the super-AGB phase thus on the metallicity (e.g., Pumo *et al.* 2009).

(2) $10 - 13 M_{\odot}$ stars: **Faint Supernovae:** These stars undergo Fe-core collapse to form a neutron star (NS) after the phase of strong Neon shell-flashes (Nomoto & Hashimoto 1988). Their Fe core is relatively small, and the resultant SNe tend to be faint (Smartt 2009).

(3) $13 M_{\odot} - M_{\text{BN}}$ stars: **Normal Supernovae:** These stars undergo Fe-core collapse to form a NS, and produce significant amount of heavy elements from α -elements and Fe-peak elements. The boundary mass between the NS and black hole (BH) formation, $M_{\text{BN}} \sim 25 M_{\odot}$, is only tentative.

(4) $M_{\text{BN}} - 80 M_{\odot}$ stars: **Hypernovae and Faint Supernovae:** These stars undergo Fe-core collapse to form a BH. SNe seem to bifurcate into two branches, Hypernovae and Faint SNe. If the BH has little angular momentum, little mass ejection would take

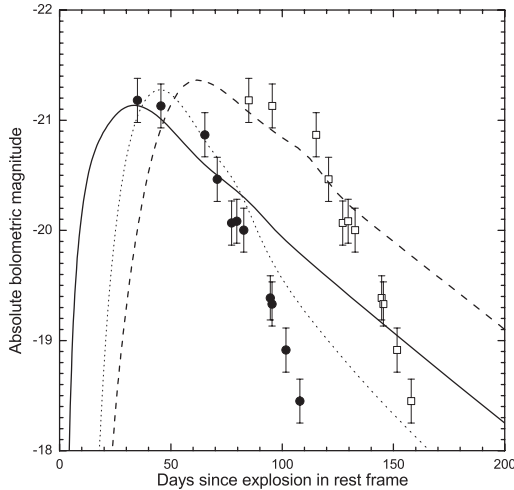


Figure 1. Model LCs with $30 M_{\odot}$ ejecta and a kinetic energy of 3×10^{52} erg, compared with the observed bolometric LC of SN 1999as, with the first data epoch of $t_0 = 35$ days (filled circles) and of $t_0 = 85$ days (open squares) are shown (Deng *et al.* 2009). The model LC to approximately fit the former was synthesized assuming $\sim 7 M_{\odot}$ ^{56}Ni (solid line), while the latter $\sim 9.5 M_{\odot}$ (dashed line).

place and be observed as Faint SNe. On the other hand, a rotating BH could eject a matter in a form of jets to make a Hypernova. The latter explosions produce a large amount of heavy elements from α -elements and Fe-peak elements. Nucleosynthesis in these jet-induced explosions is in good agreement with the abundance patterns observed in extremely metal-poor stars.

(5) **80 - 140 M_{\odot} stars: Superluminous SNe (SLSNe):** These massive stars undergo nuclear instabilities and associated pulsations (ϵ -mechanism) at various nuclear burning stages depending on the mass loss and thus metallicity. Eventually, these stars undergo Fe-core collapse. Depending on the angular momentum, Hypernova-like energetic SNe could occur to produce large amount ^{56}Ni . (Because of the large ejecta mass, the expansion velocities may not be high enough to form a broad line features.)

Thanks to the large E and ^{56}Ni mass, SNe in this mass range could be SLSNe. The upper limit of ^{56}Ni mass has been estimated to be $\sim 10M_{\odot}$. The possible presence of circumstellar matter (CSM) leads to an energetic SN II_n. Pulsation could also cause superluminous events.

(6) **140 - 300 M_{\odot} stars: SLSNe:** If these very massive stars (VMS) do not lose much mass, they become pair-instability supernovae (PISN). The star is completely disrupted without forming a BH and thus ejects a large amount of heavy elements, especially ^{56}Ni of mass up to $\sim 40M_{\odot}$. Radioactive decays could produce SLSNe (Gal-Yam in this volume).

(7) **Stars with $M \gtrsim 300M_{\odot}$: SLSNe:** These VMSs are too massive to be disrupted by PISN but undergo core collapse (CVMS), forming intermediate-mass black holes (IMBHs). Some mass ejection could be possible, associated with the possible jet-induced explosion, which could form a superluminous SNe (SLSNe).

Here we focus on the light curve models for superluminous Type Ic supernova 1999as (Deng *et al.* 2009, Nugent *et al.* 2009) to understand the explosion mechanisms of very massive stars. We find that the light curve models for SN 1999as are not quite successful and need further improvement.

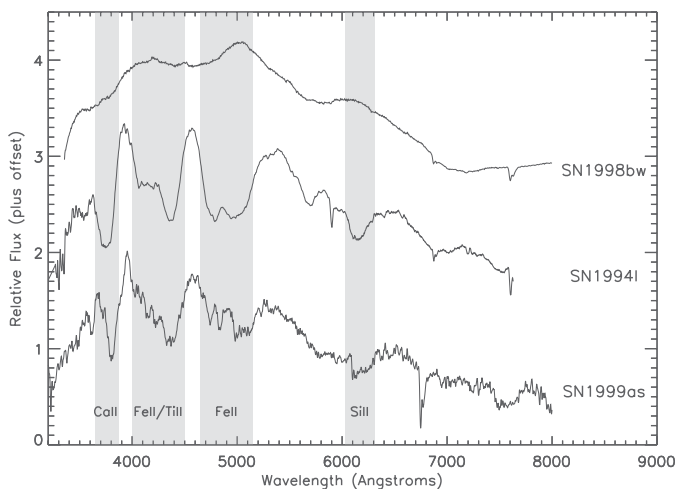


Figure 2. The spectrum of SN 1999as 25 days after discovery compared to SN 1998bw and SN 1994I (Kasen 2004).

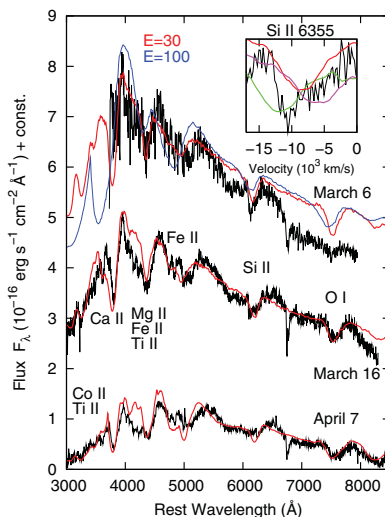


Figure 3. Synthetic spectra compared with the observed spectra taken on 1999 March 6, 16 and April 7 (from top to bottom). The red lines show synthetic spectra for the model with $30 M_{\odot}$ ejecta and a kinetic energy of 3×10^{52} ergs (Deng *et al.* 2009).

2. Superluminous Type Ic Supernova 1999as

Supernova (SN) 1999as, discovered in a galaxy of $z = 0.127$ (Knop *et al.* 1999), had long been the most luminous SN in terms of absolute brightness (Richardson *et al.* 2002, Hatano *et al.* 2001). With a peak magnitude of $M_V \sim -21.5$ (Nugent *et al.* 2009; $m - M = 38.8$), it is ~ 8 times brighter than GRB-supernova 1998bw, ~ 3 times brighter than Type Ia/IIa (IIa) SN 2002ic (e.g., Deng *et al.* 2004), and ~ 6 times brighter than normal Type Ia SNe. On the other hand, SN 1999as shows overall spectral similarities to normal Type Ic SN (SN Ic) 1994I (Fig. 2).

3. Radioactive Decay Models

The optical LC of SN 1999as (Fig. 1), like other Type Ib/Ic SNe, was possibly powered by γ -rays and positrons released in the radioactive decays of ^{56}Ni , i.e., $^{56}\text{Ni} \rightarrow ^{56}\text{Co} \rightarrow ^{56}\text{Fe}$.

3.1. Mass of Radioactive ^{56}Ni

SN 1999as had a peak $M_{\text{bol}} \sim -21.2$ (Nugent *et al.* 2009), while the spectral modeling can constrain its peak epoch to be ~ 35 – 85 days after explosion. In comparison, GRB-SN 1998bw had a peak $M_{\text{bol}} \sim -18.7$ and $t \sim 15$ days at light maximum, and its ^{56}Ni mass has been modeled to be $\sim 0.4 M_{\odot}$ (Nakamura *et al.* 2001). Then one gets $M(^{56}\text{Ni}) \sim 5$ – $9 M_{\odot}$ for SN 1999as.

From the viewpoint of explosive nucleosynthesis, such a large amount of fresh ^{56}Ni clearly indicates that the SN explosion was extremely energetic. In a core-collapse SN, most ^{56}Ni are the outcome of complete Si-burning in the region where temperatures can reach $\gtrsim 5 \times 10^9 \text{K}$ (i.e., T_{Si}) as the SN shock wave propagates through it (Arnett 1996). The region behind the shock front is radiation dominated and has roughly the same temperature (e.g., Shigeyama *et al.* 1988). Thus, complete Si-burning is confined below a radius R_{Ni} in the pre-explosion star for a given explosion energy, E , and $E \sim (4\pi/3)R_{\text{Ni}}^3 a T_{\text{Si}}^4$, or $R_{\text{Ni}} \sim 3700(E/10^{51} \text{ergs})^{1/3} \text{km}$.

As the model for SN 1999as, the explosion of the C+O core of the $M_{\text{MS}} = 80 M_{\odot}$ star leads to a final ejecta mass of $\sim 30 M_{\odot}$. As the energy rises from 1×10^{52} to 8×10^{52} ergs, $M(^{56}\text{Ni})$ increases from ~ 1 to $\sim 3 M_{\odot}$ for $M_{\text{MS}} = 50 M_{\odot}$, and from ~ 1.5 to $\sim 7 M_{\odot}$ for $M_{\text{MS}} = 80 M_{\odot}$.

For $M_{\text{MS}} = 100 M_{\odot}$ star, Type Ic SN ejecta mass is $\sim 40 M_{\odot}$. This model synthesizes $\sim 7 M_{\odot}$ ^{56}Ni for only $E \sim 4 \times 10^{52}$ ergs. The synthetic spectra will be similar for the similar E/M_{ej} .

3.2. Synthetic Spectra of SN 1999as

The spectra of SN 1999as, which were taken on 1999 March 6, 16 and April 17 (Nugent *et al.* 2009), are calculated for the $\sim 30 M_{\odot}$ ejecta produced from the explosion of the C+O core of the $80 M_{\odot}$ star with a kinetic energy of 3×10^{52} ergs.

The synthesized spectra are shown in Figure 3, and compared with the observed SN spectra. The overall fitness is satisfactory. All the absorption features are identified, being attributed to the P-Cygni lines of O I, Mg II, Si II, Ca II, Ti II, and Fe II.

The Si abundance is higher than the solar value, while Fe and Co are ~ 10 times more abundant than the solar values. This is evidence of explosive-synthesis products being mixed out into the ejecta above $8,000 \text{ km s}^{-1}$.

The supernova explosion date is also a key parameter which was estimated through spectrum modeling. The observed bolometric luminosity can be taken as approximately proportional to $(v_{\text{ph}} t)^2 T_{\text{eff}}^4$, where t is the rest-frame epoch with respect to explosion and T_{eff} the effective temperature. v_{ph} can be measured by fitting the absorption shapes of the observed P-Cygni lines (Jeffery & Branch 1990), so t is derived. A rest-frame epoch is estimated as $t_1 \sim 50$ – 100 days after explosion for the March 6 spectrum. The best-fitting model spectra discussed above correspond to $t_1 = 80$ days.

3.3. Light Curve Models

Both the LC assuming $t_0 = 35$ days after explosion and that of $t_0 = 85$ days are shown in Figure 1. The relative flatness of LC during the first month suggests that the SN was possibly discovered not far from its actual light maximum.

The bolometric LCs of $30 M_{\odot}$ ejecta with 3×10^{52} ergs are shown in Figure 1. To fit the slowest LC case, i.e., $t_0 = 85$ days, a total amount of $\sim 9.5 M_{\odot}$ ^{56}Ni is required, which

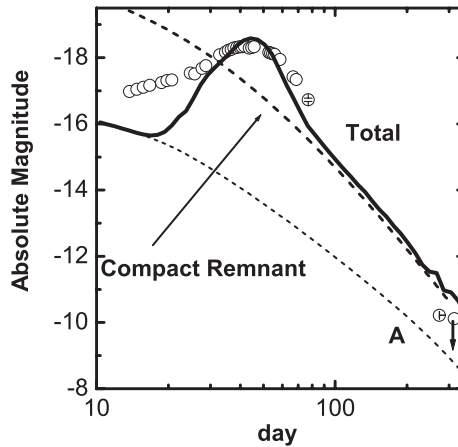


Figure 4. Synthetic light curve with the putative energy input from the central remnant (thick solid) (Maeda *et al.* 2007). The input luminosity of the remnant (thick dashed) and the contribution from the blob (thin dashed) are also shown.

were assumed by us to fill the inner part of the ejecta up to $\sim 7000 \text{ km s}^{-1}$. The LC (*dashed line*) peaks ~ 10 days before t_0 , attributable to the existence of ^{56}Ni at relatively high velocities.

3.4. Fast Decline of the Light Curve

The above radioactive decay models fail to reproduce the fast LC drop after $\sim t_0 + 60$ days. The observed luminosity decayed at a rate of $\gtrsim 0.06 \text{ mag/day}$, fast as compared with other Type Ic SNe during similar epochs ($\lesssim 0.02 \text{ mag/day}$; Tomita *et al.* 2006). In the CSM-interacting Type Ib SN 2006jc, dust formation shifted the bulk of the bolometric luminosity to the infrared (Tominaga *et al.* 2008, Nozawa *et al.* 2008). However, a pronounced red continuum should have been seen, which characterized the optical spectra of SN 2006jc at corresponding epochs (Smith, Foley, & Filippenko 2008).

4. Magnetar Models

We note that a fast late LC drop has also been observed in the peculiar Type Ib SN 2005bf, which shows two peaks in the optical light curve (Fig. 4 and references therein). The first peak luminosity is reproduced if $\sim 0.07M_{\odot}$ of ^{56}Ni is mixed out to outer layers (as a blob), and the second peak (at ~ 40 days) can be reproduced by the radioactive decay of $\sim 0.32M_{\odot}$ ^{56}Ni (Tominaga *et al.* 2005).

At ~ 270 days since the explosion, however, the absolute R-band magnitude of SN 2005bf is ~ -10.2 (Maeda *et al.* 2007). It is very faint as compared to other SNe Ib/c, at least by 2 magnitudes (e.g., by 3 magnitudes fainter than SN 1998bw at a similar epoch). If the R magnitude is close the bolometric magnitude, ^{56}Ni required to fit the luminosity is only $\sim 0.03 - 0.08M_{\odot}$. Why this is much smaller than $\sim 0.32M_{\odot}$ that reproducing the early phase peak luminosity is a serious question.

Encountered by this difficulty in the ^{56}Ni heating model, Maeda *et al.* (2007) have adopted an alternative scenario in which the heating source is a newly born, strongly magnetized neutron star (a magnetar). A synthetic light curve is shown in Figure 4 (Maeda *et al.* 2007). The input luminosity of the remnant (thick dashed) and the contribution from the ^{56}Ni blob (thin dashed) for the 1st peak are shown.

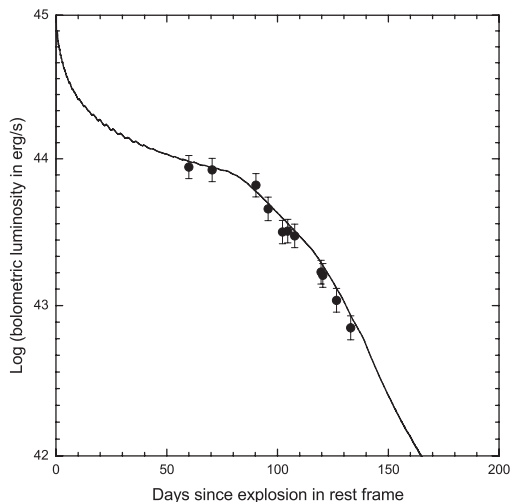


Figure 5. The LC of a CSM-interaction model (*solid line*) compared with the observed LC (*filled circles*; assuming a first data epoch of $t_0 = 60$ days). The model ejecta have $\sim 10 M_\odot$ and a kinetic energy of 1×10^{52} erg, while the model CSM have a power-law density profile index of $n = 2$ for the inner $\sim 0.90 M_\odot$ and of $n = 5$ for the outer $\sim 0.25 M_\odot$ (Deng *et al.* 2009).

The large second peak luminosity and the rapid decline after the second peak can qualitatively be explained. The model parameters are: $L_0 = 8 \times 10^{43}$ erg s $^{-1}$, $t_0 = 60^d$, $\beta = 4$, which corresponds to a pulsar with $B_{\text{mag}} \sim 3 \times 10^{14} (P_0/10 \text{ ms})^2 \sqrt{0.1/f_{\text{rad}}}$ gauss, using the dipole radiation formula. Here P_0 is the initial spin period, t_0 is the characteristic time scale of the high energy input, and f_{rad} is a fraction of energy going into the radiation. A relatively large breaking index $\beta = 4$ is required to reproduce the large contrast between the peak and the tail. The total energy injection with these parameters is $\sim 7 \times 10^{50} (f_{\text{rad}}/0.1)^{-1}$ erg, a fraction of which might be consumed to increase the kinetic energy of the SN ejecta to $E_{51} = 1 - 1.5$ and/or to develop the pulsar nebula in the early phase. For the LC of SN 1999as, it would be interesting to apply the magnetar model (also Kasen & Bildsten 2010, Woosley 2010).

5. Circumstellar Interaction Model

As seen in Figure 2, narrow spectral features are observed around 4200 and 4800 Å in SN 1999as (line width ~ 1000 km s $^{-1}$). These features seem to be high-velocity Fe lines with $v \sim 12,000$ km s $^{-1}$. We note that a jet-like explosion may result in a high-velocity shell of explosive-synthesized Fe-group elements along the jet directions (Khokhlov *et al.* 1999, Maeda *et al.* 2002). Multi-dimensional radiative transfer calculations (e.g., Höflich *et al.* 1996, Tanaka *et al.* 2007) are required to investigate such scenarios.

Kasen (2004) modeled high-velocity narrow Fe II features by assuming the existence of a thin, detached shell, and suggested circumstellar interactions for the LC of SN 1999as. Deng *et al.* (2009) calculated the circumstellar interaction model for the LC of SN 1999as in the same way as done by Suzuki & Nomoto (1995). The total ejecta mass is $\sim 10 M_\odot$. The collision starts at a distance of $r_0 = 2 \times 10^{14}$ cm, where the CSM density is ρ_0 , and the CSM has a power-law density profile $\rho = \rho_0 (r/r_0)^{-n}$.

High CSM densities and a large explosion energy are required to reproduce the brightness of SN 1999as. The density in the shocked ejecta and CSM is so high that the reverse and forward shocks are radiative to form a dense cooling shell. The regions heated by

the forward and reverse shocks emit X-rays, a large part of which is converted into UV and optical light.

The best-fitting model LC is shown in Figure 5 (*dotted line*), where the observed LC is drawn adopting an epoch of 60 days for the first data point. The model explosion energy is $\sim 1 \times 10^{52}$ ergs. The CSM parameters are $\rho_0 \sim 3.5 \times 10^{-13}$ g cm $^{-3}$, $n = 2$ for the inner $0.90 M_{\odot}$, and $n = 5$ for the outer $0.25 M_{\odot}$. The steep outer density profile is required to reproduce the late fast LC drop. The total interacting-CSM mass is then $\sim 1.15 M_{\odot}$. The corresponding mass-loss rate would be unusually large, e.g., $\sim 0.03 M_{\odot} \text{ yr}^{-1}$ within ~ 30 years prior to the SN explosion if adopting a wind velocity of ~ 100 km s $^{-1}$.

Further refinement is necessary by considering that the CSM formed with such an extensive mass loss could be optically thick. The optical thickness would improve the early part of the model LC.

6. Discussion

The optical spectra and light curve of the superluminous Type Ic SN 1999as are modeled. First we apply the radioactive model, which requires $M_{\text{ej}} \sim 30 M_{\odot}$ ($M_{\text{MS}} = 80 M_{\odot}$) and $E \sim 3 \times 10^{52}$ ergs. The high explosion energy is consistent with $\sim 6 M_{\odot}$ ^{56}Ni that are required to power the exceptional brightness of the SN. These model values of M_{ej} and E are larger than the GRB-supernovae, i.e., SN 1998bw ($\sim 10 M_{\odot}$ and $\sim 5 \times 10^{52}$ ergs; Nakamura *et al.* 2001), SN 2003dh ($\sim 7 M_{\odot}$ and $\sim 3.5 \times 10^{52}$ ergs; Deng *et al.* 2005), and SN 2003lw ($\sim 13 M_{\odot}$ and $\sim 6 \times 10^{52}$ ergs; Mazzali *et al.* 2006), although the energy value is more or less comparable.

It has been demonstrated that the core-collapse explosions of massive metal-poor stars (Umeda & Nomoto 2008) can synthesize the required large amount of ^{56}Ni of several M_{\odot} . Nugent *et al.* (2009) did find that the host galaxy of SN 1999as has a relatively low metallicity.

Pair-instability SNe from $M_{\text{MS}} \sim 140 - 270 M_{\odot}$ (Barkat, Rakavy, & Sack 1967, Bond, Arnett, & Carr 1984, Umeda & Nomoto 2002, Heger & Woosley 2002) have very large M_{ej}/E , which results in very late LC peak epochs, very slow brightness decline after the peak, and relatively narrow-lined spectra.

However, the above radioactive decay models fail to reproduce the fast decline of the light curve of SN 1999as. We thus need to explore possible alternative models.

The magnetar model has parameters to fit to the observed light curve. However, the progenitor's main-sequence mass should be smaller than $\sim 25 M_{\odot}$ to produce a magnetic neutron star rather than a black hole.

The strong circumstellar interaction models for SN 1999as would require the existence of an optically-thick, non-Hydrogen shell. Its configuration would be non-spherical to avoid the contamination of the Type Ic spectral features which dominated the observed spectra. However, the presence of non-Hydrogen CSM is not trivial. The collision of massive stars (in particular, Wolf-Rayet stars) in a dense stellar cluster can make a massive C+O star surrounded by a massive non-Hydrogen CSM (Suzuki *et al.* 2007).

The above models are not quite successful in explaining the properties of SN 1999as, in particular, the late fast LC drop and some narrow spectral features. It is crucially important to improve these models to clarify the final fates of massive stars.

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