

Time-variability and disk geometry in Herbig Ae/Be disks

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Abstract. The origin of optical-infrared variability in young, intermediate mass Herbig Ae/Be stars is linked to their circumstellar disk. Therefore, variability could serve as a diagnostic tool to constrain the structure and dynamics of the (inner) disk. Here we discuss this diagnostic potential, and report some preliminary results from our coordinated *BVR IJHK_S* and Spitzer monitoring observations of nine Herbig Ae stars. We aim to understand the response of the inner disks thermal emission on the changing stellar irradiation, and to separate it from UX Orionis-type fading events, which also provide information on the disk. This project is a pilot study for the era of time domain astronomy of young stars, opened by Kepler K2, Gaia, ASAS-SN, TESS, Spitzer, WISE, and JWST.

Keywords. stars: pre-main-sequence, circumstellar matter

1. Motivation: linking variability with disk structure

The three main physical processes causing variability are all connected to the disk: (a) fluctuating mass accretion from the disk changes the luminosity of the central source (optical light curve), and leads to variable irradiation of the disk (infrared light curve); (b) variable emitting surface of the disk, e.g. by temporary dust clouds above the disk surface, increase the thermal infrared emission; and (c) UX Ori (UXor)-type eclipses, when dust clumps in the inner disk pass through the line-of-sight. Radiative transfer models quantitatively predict the wavelength dependent light curves (Kesseli *et al.* 2016). Observationally, we looked for a correlation between the variability amplitude at 6 – 8 μm (Kóspál *et al.* 2012) and the disk flaring (Meeus *et al.* 2001). We found higher amplitudes in flatter (Group IIa) disks. Plotting the same variability amplitudes wrt. the radius of the interferometric emitting area (Menu *et al.* 2015) implies that variability in disks with more centrally concentrated brightness distribution is stronger. These results imply a physical link between variability and disk structure, and suggest that variability may provide information about the disk structure and dynamics on the smallest spatial scales, a few au from the star.

2. The Konkolyvar project

In order to decide which physical processes are responsible for the observed variability in a particular system, and to deduce disk properties, simultaneous optical-infrared observations are needed. We carried out a 2-week-long quasi-simultaneous optical-infrared photometric monitoring program with daily cadence, in 2009 Oct-Nov. We targeted nine

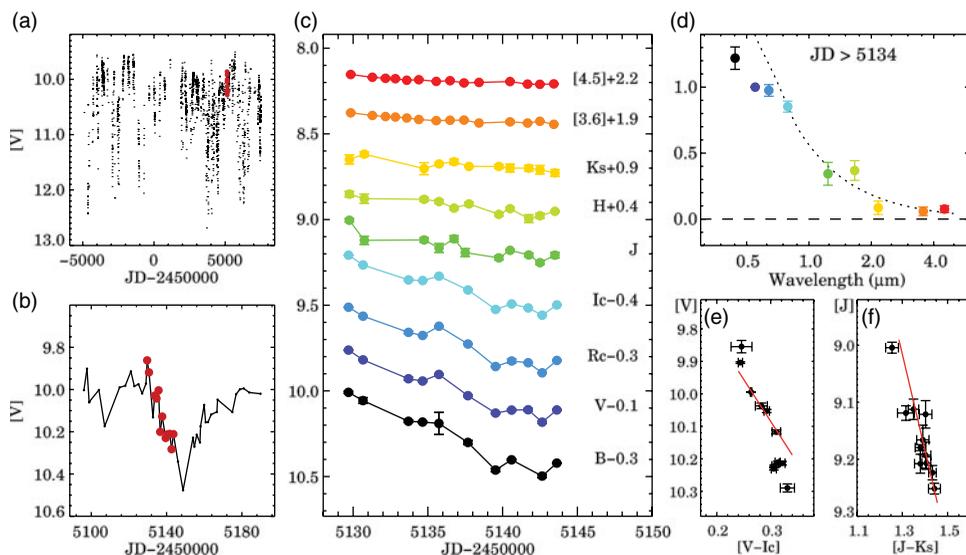


Figure 1. Multiwavelength light curves for BF Ori from the Konkolyvar project.

Herbig Ae stars and collected observations from Pizskéstető Observatory (Hungary), Teide Observatory (Spain), and at 3.6 and 4.5 μm from the Spitzer Space Telescope.

3. The physical origin of the variability, and the response of the disk

As an example, Fig. 1 presents our observations for BF Ori, where panel *c* shows the multiwavelength light curves. Comparison with earlier/later *V*-band observations (panel *b*) outlines that our first two epochs coincided with a small brightness peak, while later the source entered a deep minimum. Panels *d*, *e*, and *f* imply that the fading occurred along the extinction path, and was probably a typical UXor event. The faintest data points exhibit optical colors deviating from the extinction curve: here we see the well-known blueing of UXors, caused by scatter light.

We also compared the Spitzer 3.6/4.5 μm data with the *I*-band light curve. If the whole inner disk was obscured in an UXor event, shapes of the *I*-band and Spitzer light curves would be identical, apart from a lower amplitude in the mid-infrared according to the extinction law (e.g., Cardelli *et al.* 1989). This is the case for BF Ori at JD > 245135, but the first few points exhibit the response of the disk with unexpectedly high amplitude for the small *V*-band brightness peak. Analysing the other stars from the Konkolyvar sample we found large variety: in RR Tau most variability is related to extinction changes, while in UX Ori the disk emission is virtually unrelated to the optical flux changes, because in the second part of our campaign the optical flux increases while the mid-infrared flux drops. The detailed analysis of these results in terms of disk geometry and dynamics is ongoing.

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