

Morphological Structures of Planetary Nebulae

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Abstract: Since various structural components of planetary nebulae (PN) manifest themselves differently, a combination of optical, infrared, submillimetre and radio techniques is needed to derive a complete picture of planetary nebulae. The effects of projection can also make derivation of the true 3-D structure difficult. Using a number of examples, we show that bipolar and multipolar nebulae are much more common than usually inferred from morphological classifications of apparent structures of planetary nebulae. We put forward a new hypothesis that the bipolar and multipolar lobes of PN are not regions of high-density ejected matter, but the result of ionization and illumination. The visible bright regions are in fact volumes of low density (cleared by high-velocity outflows), through which UV photons are being channelled. We suggest that multipolar nebulae with similar lobe sizes are not caused by simultaneous ejection of matter in several directions, but by leakage of UV photons in those directions.

Keywords: planetary nebulae: general — stars: AGB and post-AGB — stars: mass loss

1 Introduction

Planetary nebulae (PNe) are traditionally identified by their morphological appearances in photographic and imaging surveys, or by their spectroscopic properties in objective prism or emission-line surveys. In recent years, new PNe candidates have also identified through their spectral properties among objects in radio surveys, or through their colors among objects in infrared surveys. Although a number of observational properties have been employed as selection criteria for PNe, extensive confusion still exist between other classes of objects such as HII regions, symbiotic stars, Wolf–Rayet nebulae (Kwok 2000).

Recent narrow-band imaging CCD surveys such as the Anglo–Australian Observatory, UK Schmidt Telescope H α Survey (Parker et al. 2005; Miszalski et al. 2008) and the Isaac Newton Telescope Photometric H α Survey of the Northern Galactic Plane (IPHAS) (Drew et al. 2005) have greatly expanded the number of PNe candidates. Comparison between the brightness in the narrow-band (H α) and broad-band images easily identifies the emission-line objects, and the images allow for morphological identification. The PN nature of the candidates can be further strengthened by follow-up spectroscopic observations (Parker et al. 2006). The expanded list of PNe in the Macquarie/AAO/Strasbourg H α Planetary Nebula Catalogue (MASH) allows for a more systematic study of the morphology of PNe. For example, the improved sensitivity of the recent surveys makes possible the detection of low-surface brightness objects, which could represent a population of evolved PNe.

2 Morphological Classification

PNe are classified into morphological classes based on their apparent structure. Beginning with Curtis (1918), there have been many attempts in morphological classifications based on observational surveys of PNe (e.g. Khromov & Kohoutek 1968; Stanghellini, Corradi & Schwarz 1993; Aaquist & Kwok 1996; Manchado et al. 1996). Parker et al. (2006) classify the 903 new PNe in the MASH catalogue into 6 classes: round, elliptical, bipolar, irregular, asymmetric and star-like. The most basic classification is that of Balick (1987), who classifies PNe into classes of round, elliptical and butterfly. The morphological classification studies have often emphasized the small fraction ($\sim 15\%$) of PNe having bipolar structures. Is this small percentage an accurate reflection of the fraction of PNe with bipolar morphologies? Often the true intrinsic structures of PNe cannot be easily derived from the observed images, as all classification schemes suffer from the following problems:

• Sensitivity Dependence

A deeper exposure can reveal fainter structures which change the classification of the PNe. For example, the waist of a bipolar nebula could be classified as elliptical if the bipolar lobes are too faint to be detected. NGC 650-1, Sh 1-89 and SaWe 3 are some of the cases where their bipolar nature were only revealed after deep CCD imaging (Hua 1997; Hua, Dopita & Martinis 1998).

• Species Dependence

The morphology of PNe observed in lines of different ions is not necessarily the same, as the result of ionization structures and stratification effects.

• Projection Effects

Morphology classifications describe the two-dimensional apparent structures, not the intrinsic structures of the PNe (Khromov & Kohoutek 1968; Manchado 2004).

A bipolar nebula with an equatorial waist viewed at an angle other than edge on will have a ring-like appearance, especially when the bipolar lobes are faint. The best-studied PN NGC 7027 has a very prominent ring structure, but when studied in detailed, it is revealed to be a bipolar nebula (Latter et al. 2000). The same is true for NGC 3132 (Monteiro et al. 2000). Even well-known objects such as NGC 6720 (the Ring Nebula) and NGC 7293 (the Helix Nebula) turn out to have bipolar morphologies (Bryce, Balick & Meaburn 1994; Meaburn et al. 2005a). Two objects with very different apparent morphologies, such as the Ring and NGC 6853 (the Dumbbell Nebula), may in fact have similar intrinsic 3-D structures (Kwok et al. 2008).

In Figures 1 and 2, we show models of MyCn 18 and NGC 7009, respectively. By rotating the model images to another orientation, we can see that same objects could resemble other well known PNe, namely NGC 7293 (the Helix Nebula) and NGC 6826, respectively. A similar 3-D reconstruction exercise has also been done by Sabbadin et al. (2004). These examples clearly illustrate that apparent morphology alone is not sufficient to obtain the true intrinsic structure of PN. Kinematic data are necessary to separate various components projected on the same positions in the sky.

3 Multi-Wavelength Observations of Planetary Nebulae

A common perception of PN is that a shell of gas photo-ionized by a hot central star. Now we know that a typical PN is made up of ionized, atomic, molecular and dust components, and that these components do not necessarily have the same geometry. Their temperatures range from 100 to 10^6 K and need to be studied with different observational techniques from radio to X-ray. The optical component, the most easily observed and therefore the best studied, show multiple structures in the form of shell, crown, halo and lobes. These structures have been successfully modeled by 1-D dynamical evolution of PNe incorporating stellar evolution, interacting winds dynamics and time-dependent photoionization (Perinotto et al. 2004; Steffen & Schönberner 2006).

The existence of a dust torus is often inferred from the presence of a dark lane in the optical image of bipolar nebulae. Mid-infrared imaging can provide a direct determination of the size and orientation of the torus (Volk et al. 2006). In the proto-PN IRAS 17441–2411 where the orientation of the axis of the infrared torus is found to be offset by 23° from the bipolar axis. If the torus is responsible for or related to the collimation of the bipolar outflow, this raises the possibility that the outflow has undergone precession over the last 100 years (Volk, Kwok & Hrivnak

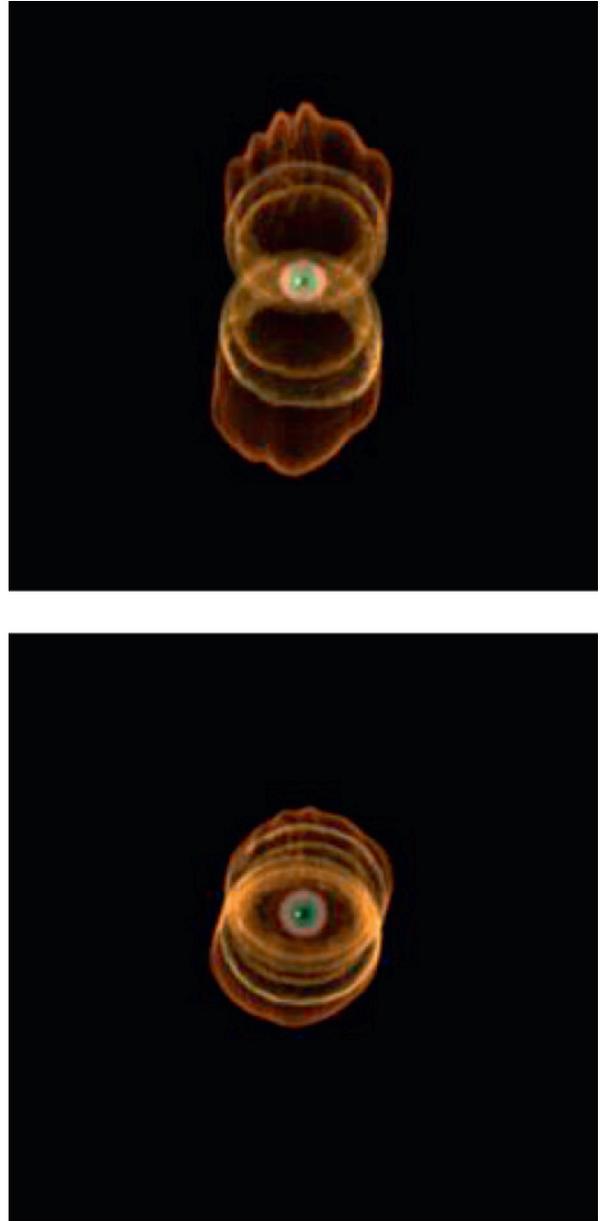


Figure 1 A SHAPE (Steffen & Lopez 2006) model image of MyCn 18 (top) and the same model rotated to another viewing angle that resembles the Helix Nebula (bottom). The model was first created to animate the *HST* image of MyCn 18 (Sahai et al. 1999) and then rotated to other orientations. Models in Figures 1 and 2 are constructed by N. Koning.

2007). From observations obtained with the Multiband Imaging Photometer (MIPS) of the *Spitzer Space Telescope*, Su et al. (2004) show that warm dust (as evidenced by emission at $24 \mu\text{m}$) is present in both the bipolar lobes and in the torus of NGC 2346. Cool dust (as evidenced by emission at $70 \mu\text{m}$) is mostly located in the edge-on equatorial torus and the cold dust (as evidenced by emission at $160 \mu\text{m}$) is distributed over a spherical region, probably arising from the remnant of the AGB wind.

With sub-mm arrays, such as *PdBI*, *CARMA*, *SMA* and *ALMA*, the distribution of the emission regions of molecular lines, free-free continuum and dust continuum can be mapped with sub-arcsec angular resolutions. Comparisons

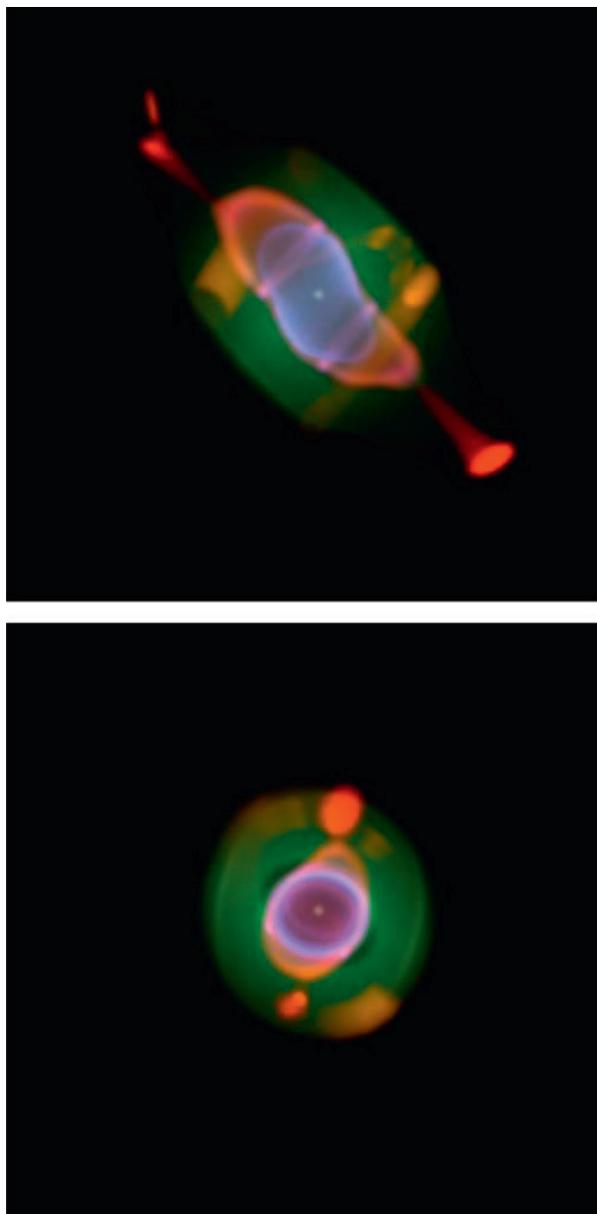


Figure 2 A SHAPE model image of NGC 7009 (top) and the same model rotated to another viewing angle that resembles NGC 6826 (bottom). The model was first created to animate the *HST* image of NGC 7009 (Balick et al. 1998) and then rotated to other orientations.

between the molecular, optical and radio continuum maps of NGC 7027 show a clear photodissociation region separating the ionized and molecular gas components (Graham et al. 1993). For NGC 6302, sub-mm observations have found an expanding molecular torus, which correspond to the dark lane separating the optical bipolar lobes (Peretto et al. 2007; Dinh et al. 2008).

Fast molecular outflows can be detected by single-dish mm/sub-mm observations from the extended wings of the CO profile. High-resolution CO $J = 6-5$ images of AFGL 618 show that the high-velocity components lie along the optical lobes (Nakashima et al. 2007). In NGC 2440, the CO outflow has been imaged to correspond to one pair of the optical bipolar lobes (Wang, Hasegawa & Kwok 2008).

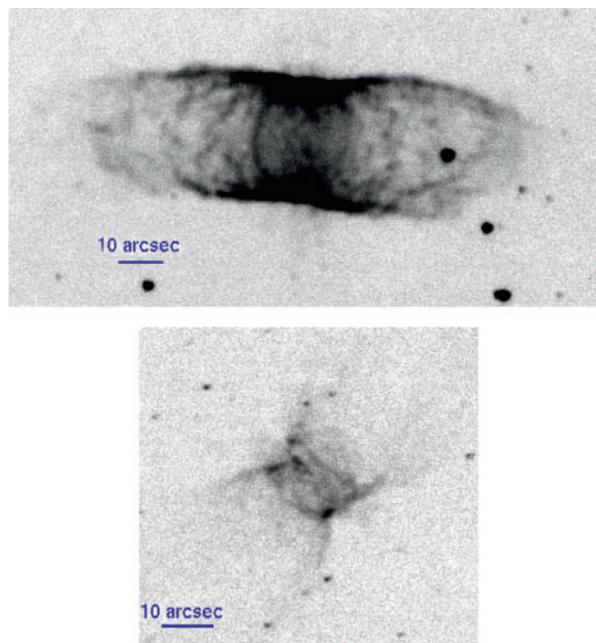


Figure 3 Continuum-subtracted H_2 image of IC 4406 (top) and K3-72 (bottom) obtained with the Canada-France-Hawaii Telescope. The bipolar nebulae, especially the waist of the nebulae, are clearly confined externally by unseen neutral gas.

With the construction of *ALMA*, we will be able to map the dust and molecular components of PNe with much greater sensitivity and angular resolution. Only through a comprehensive study of the ionized gas, molecular gas and dust components, can we arrive at a full picture of the morphology of PNe.

4 Molecular Hydrogen as a Tracer of Dynamical Interactions

The 2.12- μm vibrational-rotational line of H_2 is a useful tracer of dynamical interactions in PNe. A narrow-band H_2 image can often reveal structures not seen in broad-band images. Figure 3 shows H_2 image of two bipolar PNe IC 4406 and K3-72. In both cases, rays of equatorial outflows can be seen. Since the IC 4406 is inclined at an angle, a bright elliptical ring can clearly be seen in the waist region. The bright rims in the lobes of IC 4406 sharply define the edges of the lobes, suggesting that the lobes are confined by an external medium and that the H_2 emission trace regions of wind interactions. The H_2 emissions at the caps of the lobes probably represent the regions where the fast wind is breaking through the remnant AGB circumstellar envelope. We note that these features are much more prominent in this H_2 image than in the *HST* optical image of IC 4406 (O'Dell et al. 2002). The equatorial features resemble the high-speed 'skirts' seen around the waist of NGC 6302 (Meaburn et al. 2005b), or the 'chakram' of Mz-3 (Santander-Garcia et al. 2004). Although K3-72 has a more edge-on orientation, similar structures of equatorial torus, bipolar lobes and radial equatorial jets can also be seen.

We can infer from Figure 3 that the tight waists of IC 4406 and K3-72 are due to confinement by an unseen

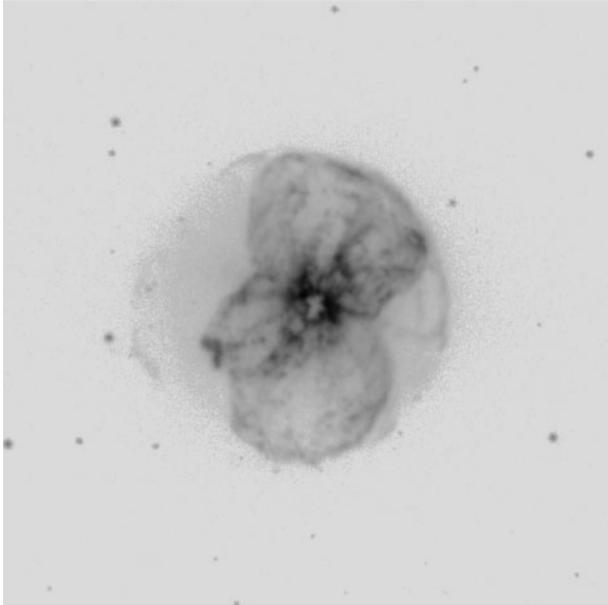


Figure 4 A [N II] image of NGC 2440 taken at the Canada-France-Hawaii Telescope showing two pairs of bipolar lobes. The round outer shell probably represent the extent of the AGB envelope. Are the multiple lobes manifestations of separate ejection events or photoionized regions bounded by neutral gas? (see discussions in Section 6).

equatorial torus and the nebulae are ionization bounded in the equatorial directions. Infrared and sub-mm imaging will be needed to determine the geometric and dynamical relationship between the torus and the bipolar lobes.

5 Multipolar Nebulae

Although planetary nebulae are commonly believed to possess axi-symmetry, an increasing number of planetary nebulae has been found to be point-symmetric, including the class of multipolar nebulae (Manchado, Stanghellini & Guerrero 1996). The most well-known example is NGC 2440, which in addition to its bright pair of bipolar lobes, have been found to have two other fainter pairs of lobes (López et al. 1998, see also Figure 4). A survey of compact planetary nebulae by the *HST* has also revealed further multipolar objects such as M1-37 and He2-47 (Sahai 2000) and NGC 6881 (Kwok & Su 2005). NGC 6072, a bipolar PN with a bright waist (Corradi & Schwarz 1995), in fact possess two or three pairs of bipolar lobes with the bright central part being an equatorial ring (Figure 5).

At the present time, we may be just scratching the surface of the multipolar phenomenon. Mampaso et al. (2006) have found a new PN with an equatorial ring, a pair of inner lobes, a pair of bright main lobes and very large faint outer lobes. The symmetry axes of these lobes all seem to have different orientations. Further deep and wide-field imaging of PNe will lead to a greater appreciation of the complexity of the PN morphology.

If we model a multipolar nebula (e.g. by systems of equatorial rings plus bipolar lobes) and view the objects from different perspectives, we get a very different impression of the object. For example, the equatorial rings viewed

near pole on will be taken as projection of a spherical shells, and the projections of the other lobes will be assumed to be internal structures of the shell. The apparent morphologies of these different projections can easily mislead the observer about the intrinsic structure of the nebula.

6 Effects of Illumination

It is important to remember that an optical image of a PN does not give a complete representation of the distribution of matter. We have already mentioned that there is ‘dark matter’ in the form of molecular and solid-state matter in PNe and that their masses can be considerably higher than the total mass of the ionized region. I have long argued that since PPN shine by reflected star light, the visible image that we see in a PPN in fact represents the cavity where light can escape, completely opposite to the actual distribution of matter (Kwok 2004). Even in PNe, the bright optical emission region may represent low-density regions cleared out by high-velocity outflows. In bipolar nebulae such as IC 4406 shown in Figure 3, the optical nebula is probably ionization bounded and bipolar lobes are ionized because these are the directions where the UV photons are channeled through.

In this model, the morphology of PNe is not so much defined by regions of matter ejection, but are defined by ‘holes’ in the matter distribution where densities are low enough for the available UV flux to ionize. We cannot have a complete description of matter distribution until we have mappings of both the ionized and neutral gas components.

7 Dynamical Evolution

The fraction of PNe with bipolar structure has important implications on our understanding of PNe dynamical evolution. If the real fraction of bipolar PNe is indeed high, then one is compelled to seek a physical scenario of how such a morphological transformation occurs. From molecular-line imaging, we know that the mass-loss process is mainly spherically symmetric, with the possible departure from spherical symmetric at the very end. In the following post-AGB evolution, observations of proto-planetary nebulae suggest that one or more highly collimated outflows develop, carving out bipolar cavities in the spherical remnant AGB envelope. Shell structures develop as the result of the sweeping up of remnant AGB material by these fast outflows. Although the above scenario roughly describes the morphological transformation of PNe, the details are not known. For example, it is not clear whether an equatorial outflow develop first, forming an equatorial disk which collimates the bipolar flow, or vice versa (Huggins 2007).

An interesting aspect of the multi-polar phenomenon is that the sizes of the multi-polar lobes are approximately the same, suggesting that they originate at approximately the same time. How a PN can simultaneously eject several pairs of highly collimated outflows is a mystery. Most models attribute the multipolar phenomenon to precession of the collimating disk, with either binarity or magnetic as the agent. Here we offer an alternative explanation.

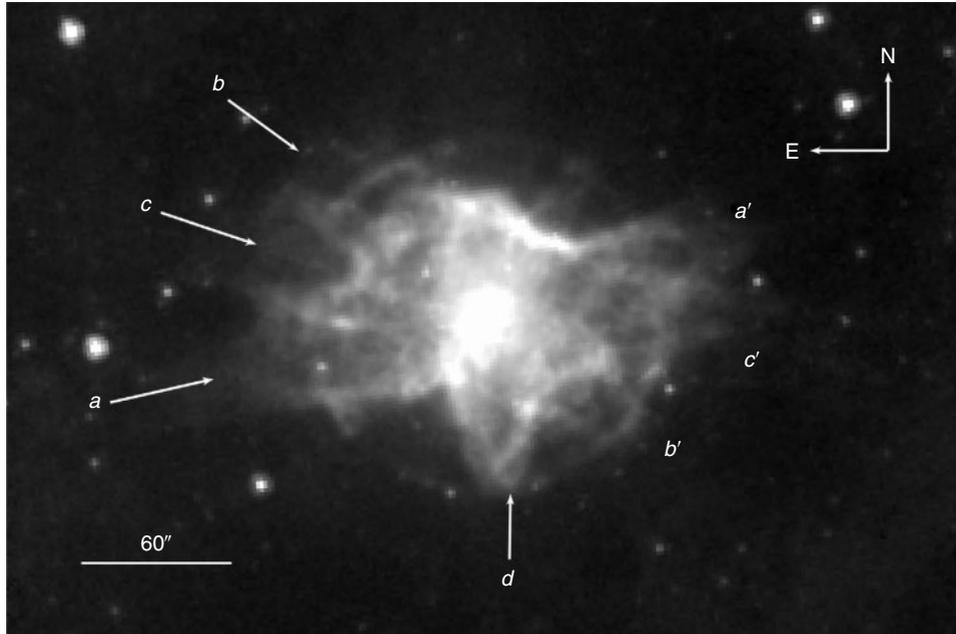


Figure 5 Spitzer IRAC 8- μm image of NGC 6072 with the three possible pairs of bipolar lobes marked as a – a' , b – b' , c – c' . The equatorial disk is marked as d . Are the multiple lobes regions of separate ejections or directions of UV photon leakage?

According to the theory outlined in Section 6, the multipolar lobes only reflect directions where the equatorial torus has ‘holes’. If this is the case, the explanation to the multipolar phenomenon lies not so much in the dynamical ejection, but in how such holes are created. The similarity in sizes of the multiple lobes is therefore naturally explained and no simultaneous ejection is required.

The MASH survey has found many low-surface brightness, round, PNe. If they correspond to an evolved stage of evolution of PNe, one has to account for the dynamical morphological evolution from bipolar back to spherical symmetry. This, e.g., can be attributed to the increasing dominance of the effect of thermal pressure of the bubble and the effect of recombination as the UV output of the central star declines (Steffen & Schönberner 2006). If we again appeal to the theory put forward in Section 6, an evolved PN will be entirely density bounded and the spherical morphology just reflects the original morphology of the AGB envelope.

8 What is a Planetary Nebula?

Planetary nebulae are usually defined by a combination of observation properties of the nebula by morphology (having some degree of symmetry), by spectrum (emission-line spectrum with little or no continuum, $H\alpha$ to $[\text{O III}]$ ratio, expansion velocity, etc.) and by the properties of the central star (temperature, luminosity, gravity, etc.). However, many symbiotic stars (in particular those of D-type) share similar properties. Symbiotic stars and PNe even have similar properties in the radio, infrared and X-ray spectral regions. There are some distinctions, however. Symbiotic stars have higher excitation lines as the result of accretion, periodic photometric variability due to the pulsation of the cool component and molecular

absorption features produced in the atmosphere of the cool star (Kwok 2003). It has also been suggested that symbiotic stars have more extreme bipolar morphologies, e.g. as in Hen 2-104 and R Aqr (Corradi 2003).

When a new class of astronomical object is first discovered, they are usually first defined by their observational properties. But as we gain insight into the phenomenon, a more precise definition is given based on the physical or evolutionary nature of the class of objects. For PNe, we can define them as ionized expanding circumstellar shells showing some degree of symmetry ejected by a hot, compact, central star evolving from the AGB to the white dwarf phase (Kwok 2000). This definition will distinguish PNe from symbiotic stars or novae, which are binary systems undergoing mass exchange. While the hydrogen envelope mass of the a PN central star is being depleted by nuclear burning and mass loss, and therefore constantly evolving, a symbiotic star maintains its energy source through accretion and is therefore stationary (or even going backwards) in evolution (Paczynski & Rudak 1980). It should be noted, however, that a post-outburst nova or symbiotic nova also evolves to the blue similar to a PN. Unless the outburst (H ignition) is observed, it may also be difficult to distinguish a symbiotic nova (e.g. V1016 Cyg and HM Sge) from a PN. The separation of symbiotic stars from PNe in PN catalogues will be a long continuous process (Kohoutek 1994). For PNe in external galaxies, this is almost an impossible task.

Since a binary system can undergo mass exchange either when the primary or the secondary is in the post-main-sequence phase of evolution, the number of evolution scenarios (depending on initial mass ratios and separations) is almost infinite (Webbink 1988). When a binary system is observed in a PN, relevance to PN

evolution can be none (for very wide binaries) to a complicated history, with the symbiotic phenomenon being one of the possible outcomes.

9 Conclusions

Recent observations have revealed PN structures that are much more complicated than a spherical shell. Even a seemingly simple symmetric nebula such as the Ring Nebula in fact has a multiple bipolar outflow history. While we have learned that interacting winds plus time-dependent photoionization as the result of stellar evolution can lead to multiple-shell structures (Schönberner et al. 2005), they can also be the result of projection of bipolar or multipolar lobes plus torus. Such confusion can only be sorted by kinematics. The traditional long-slit spectroscopy does provide a certain amount of information, but we really need complete kinematic information of the entire nebula, including the shell, the crown, the lobes and the halo. For the ionized region, this can be achieved by integral field spectroscopy and for the neutral envelope, by molecular line imaging. Together with 3-D modeling (e.g. *SHAPE*), we may be able to derive the true 3-D structure.

An understanding of the real geometric structure of PNe is the first step toward understanding of their dynamical evolution. This should begin with wide-field, high dynamic-range, multi-wavelength, imaging spectroscopic studies of the brightest PNe. Only after we have gained a complete understanding of the 3-D structures of PNe such as the Ring, the Dumbbell, the Helix, would we be in a position to talk about the structure of other fainter PNe with confidence.

Our common perception of morphology of PNe is guided by their optical appearances. Since ionized gas represents a small fraction of the total mass of PNe, optical structures cannot be used to infer the total mass distribution. Regarding the physical cause of the bipolar/multipolar phenomena, we raise the possibility that this is not so much the result of dynamical ejection, but effects of illumination and ionization. The bipolar lobes are in fact ionization-bounded low-density regions confined by high-density neutral gas. This model can be tested by high angular resolution imaging of the molecular and dust components, a task that can be accomplished by *ALMA*.

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