MILLIMETER-WAVE MAPPING OF THE W3 CORE REGION AT 4 AND 6.5 mm

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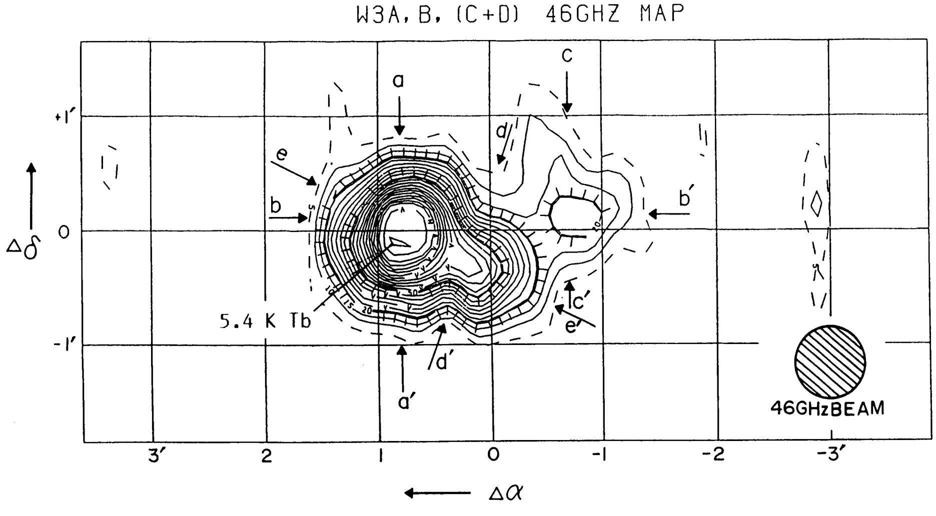
Radio continuum observations of the core region of the compact H II region W3 at 6.5 and 4 mm wavelengths were made using the Nobeyama 45-m telescope. The 6.5-mm map agrees with lower-frequency maps, showing a major contribution of free-free H II emission. At 4 mm an excess over the H II emission is found, which indicates a contribution from dust grains. Comparing with sub-mm and FIR data, we suggest the existence of two dust components: normal dust at 50 K, and low-temperature (7 K), large-size grains (or interstellar "stones") in the region west of the W3 core.

The 46-GHz map (Figure 1) has a simple structure convolved from the three point-like sources W3A, B and C+D and a weak extended component north of C+D. The 75-GHz map (Figure 2) looks similar to the 46-GHz map. However, at 75-GHz a more extended component than at 46-GHz is found near W3-C+D. The 75-GHz excess relative to the 46-GHz emission is shown in Figure 3. The excess is distributed in the SW edge of the W3 core and is strong in the C+D region. The 46-GHz emission may mostly come from optically thin H II gas, while the 75-GHz emission comes from both the H II gas and dust grains.



(0,0): (α= 02 h21m50s, δ= +61°52'40". 1950)

Fig. 1. A 46-GHz map of W3. The W3 peak contour of 120 corresponds to 5.4 K tb.

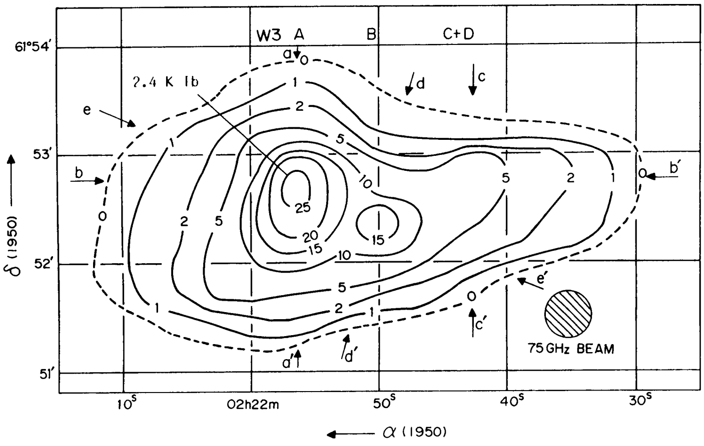


Fig. 2. A 75-GHz map of W3. The peak contour of 25 corresponds to 2.4 K Tb.

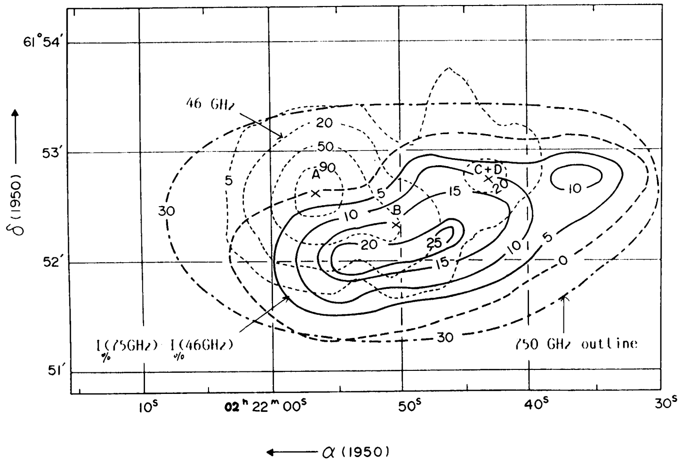


Fig. 3. Excess in the 75-GHz emission over the 46-GHz emission. Contour levels are in percentage of the W3 peak.

A spectrum of the excess brightness over the free-free emission is shown in Figure 4. The spectrum at $\lambda < 1$ mm may be fitted with an optically thin warm dust of 50 K with an optical depth depending on ν as $\tau \propto \nu^2$ and $\tau = 0.02$ at 750-GHz. However, at $\lambda > 1$ mm we have a difficulty to fit the spectrum with this dust component alone. We may therefore introduce a second component so that the observed brightness can be ex-

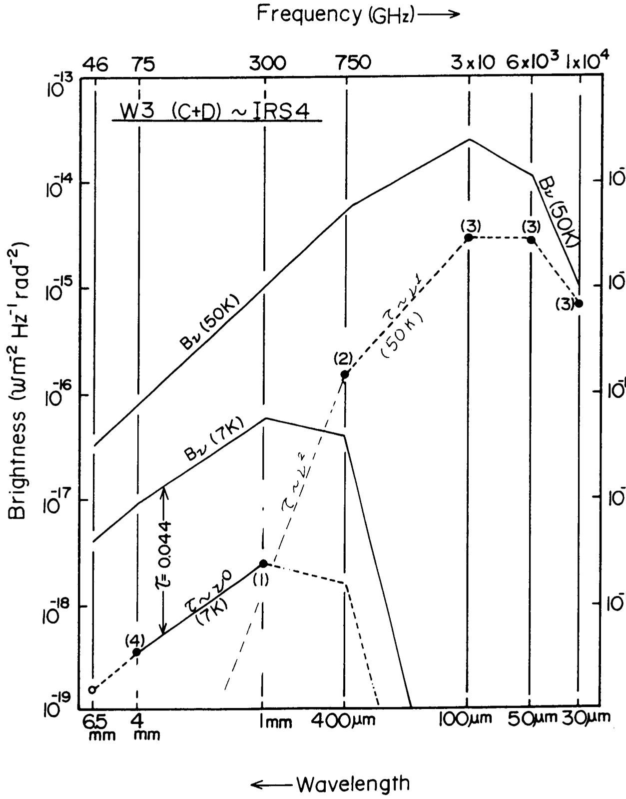


Fig. 4. Intensity spectrum on W3 C+D. The free-free H II emission has been subtracted. Plotted data are: (1) Westbrook *et al.* (1976); (2) Jaffe *et al.* (1983); (3) Werner *et al.* (1980).

pressed as $B(\nu) = \tau_1(\nu)\ B_\nu(T_1) + \tau_2(\nu)\ B_\nu(T_2)$, where τ_i is the optical depth as a function of the frequency and $B_\nu(T)$ represents the Planck's function with T_i the dust temperature. For the first term we have T_1 = 50 K and $\tau_1(\nu) = 0.02\ (\nu/750\ \mathrm{GHz})^2$. The second component can be fitted with T_2 = 7 K and $\tau_2 = 0.044$.

The total luminosity of the 50-K component is L(50 K) $\sim 10^5\ L_\odot$ and the total grain mass is several $M_\odot$; the total gas and dust in the W3 core region is $10^3\ M_\odot$ for a dust-to-gas ratio of 10^{-2} (Jaffe *et al.* 1983). This component is like a typical dust cloud normally found in compact H II regions (Schwartz 1982).

The characteristics of the second 7-K component are unclear. As the frequency dependence is small, $\tau \propto \nu^0$, we may suppose that the particle size is large compared with λ, or a >6 mm. This implies that the material is possibly "stones" rather than grains. The total luminosity of this component is L(7 K) $\sim 4\ L_\odot$. If we assume that the particle radius is $a \sim 1$ cm and density $\rho \sim 1\ \mathrm{g\ cm^{-3}}$, the total mass of the stones is $M(7\ \mathrm{K}) \sim 4\pi/3\ \rho a^2\ \ L/4\pi a^2 \lambda T^4\ \sim 10\ M_\odot$ with λ the Stefan-Boltzmann constant. It must be noticed that the "stone-to-gas" ratio is almost equal to the dust-to-gas ratio. It seems urgent to develop the physics for such a new species of interstellar material.

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THE STRUCTURE OF THE W49A MOLECULAR CLOUD COMPLEX: BURST OF STAR FORMATION IN THE 10^5 $M_\odot$ CORE

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We have observed the CS (J = 1-0), $C^{34}S$ (J = 1-0) and H51α emission toward the W49A molecular cloud complex in an area of 3'× 2' (α × δ) with an angular resolution of 33". The CS emitting region is 100" × 80" or 6.7 pc × 5.4 pc (α × δ) at the half maximum level. Although the CO emission is self-absorbed due to the foreground cold gas, the CS optical depth of the foreground gas is found to be small. Therefore, the two CS peaks at V_{LSR} = 4 km s^{-1} and 12 km s^{-1} imply the presence of two dense molecular clouds toward W49A. The brighter 12 km s^{-1} cloud peaks 35" southeast of W49A IRS, the infrared and H_2O/OH maser sources associated with the compact H II region, while the 4 km s^{-1} cloud has a peak at W49A IRS. The hydrogen column density through the $C^{34}S$ emitting region is $(0.3-1.7) \times 10^{24}$ cm^{-2}. The estimated core mass of the W49A molecular cloud is $(0.5-2.5) \times 10^4$ $M_\odot$. This mass is closely packed in a small region of 3.4 pc in diameter, and is about an order of magnitude larger than the virial mass of the system. The massive core will collapse within 10^5 years unless there is some special supporting mechanism. There was a sudden increase in the star formation rate 10^4- 10^5 years ago, suggesting a triggered burst of star formation in the core of W49A. The collision